

A high black hole to host mass ratio in a lensed AGN in the early Universe

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Early JWST observations have uncovered a new population of red sources that might represent a previously overlooked phase of supermassive black hole growth [1–3]. One of the most intriguing examples is an extremely red, point-like object that was found to be triply-imaged by the strong lensing (SL) cluster Abell 2744 [4]. Here we present deep JWST/NIRSpec observations of this object, Abell2744-QSO1. The spectroscopy confirms that the three images are of the same object, and that it is a highly reddened ($A_V \simeq 3$) broad emission line Active Galactic Nucleus (AGN) at a redshift of $z_{\text{spec}} = 7.0451 \pm 0.0005$. From the width of $\text{H}\beta$ ($\text{FWHM} = 2800 \pm 250 \frac{\text{km}}{\text{s}}$) we derive a black hole mass of $M_{\text{BH}} = 4_{-1}^{+2} \times 10^7 M_{\odot}$. We infer a very high ratio of black hole to galaxy mass of at least 3 %, an order of magnitude more than is seen in local galaxies [5], and possibly as high as 100 %. The lack of strong metal lines in the spectrum together with the high bolometric luminosity ($L_{\text{bol}} = (1.1 \pm 0.3) \times 10^{45} \frac{\text{erg}}{\text{s}}$) indicate that we are seeing the black hole in a phase of rapid growth, accreting at 30 % of the Eddington limit. The rapid growth and high black hole to galaxy mass ratio of A2744-QSO1 suggest that it may represent the missing link between black hole seeds [6] and the first luminous quasars [7].

A2744-QSO1 [4] was discovered in JWST/*Near Infrared Camera* [8] (NIRCam) pre-imaging from the *Ultra-deep NIRSpec and NIRCam Observations before the Epoch of Reionization* [9] (UNCOVER) program, behind the *Hubble Frontier Fields* [10] (HFF) SL cluster Abell 2744 [11] ($z = 0.308$; A2744 hereafter), known to be capable of significantly magnifying distant background objects [12]. It had been identified as a UV-faint ($M_{\text{UV},1450} = -16.98 \pm 0.09$), lensed high-redshift candidate [13], but remained an otherwise inconspicuous source in previous $0.5 - 1.5 \mu\text{m}$ observations with the *Hubble Space Telescope* (HST). The NIRCam $1 - 5 \mu\text{m}$ imaging revealed it to be an unusually red (e.g. $\text{F277W} - \text{F444W} = 2.63 \pm 0.10$), unresolved point-source triply imaged by the foreground cluster [4], with the high-redshift nature of the source further supported geometrically by the lensing configuration. Its unique HST and JWST spectral energy distribution (SED), combined with the point-like light-profile, hinted that the source may be a highly reddened AGN at redshift $z_{\text{phot}} \simeq 7 - 8$ [3, 4].

The three images of A2744-QSO1 were recently targeted with the *Near Infrared Spectrograph* [14] (NIRSpec) aboard JWST as part of the UNCOVER program (see Fig. 1). Stacking the spectra of the three images of A2744-QSO1, we obtain one ~ 38 h spectrum, the deepest spectrum so far of any single high-redshift object observed with JWST. After correcting for gravitational magnifications [12] of $\mu_{\text{A}} = 6.2_{-0.2}^{+0.5}$, $\mu_{\text{B}} = 7.3_{-1.5}^{+0.1}$ and $\mu_{\text{C}} = 3.5_{-0.2}^{+0.2}$, this corresponds to a 1.5 micron continuum depth of ~ 31.0 AB at 5σ , equivalent to integrating for $\gtrsim 1700$ h in a blank field.

The spectrum shown in Fig. 2 features several strong Hydrogen emission lines – a prominent Balmer-series with $\text{H}\alpha$, $\text{H}\beta$, $\text{H}\gamma$, and $\text{H}\delta$, as well as Lyman- α – along with several faint metal lines. We give a full list of detected emission lines in Tab. 1. The broad Hydrogen emission lines securely confirm this object at $z_{\text{spec}} = 7.0451 \pm 0.0005$ and provide unambiguous evidence that A2744-QSO1 is an AGN. The broad-line width is measured by fitting the $\text{H}\beta$ line, which is the clearest and most isolated

since $H\alpha$ is situated on the edge of the detector. We thus find an $H\beta$ full width at half-maximum (FWHM) of $\text{FWHM} = 2800 \pm 250 \frac{\text{km}}{\text{s}}$, well-resolved, even at the relatively low resolution of the prism (see Fig. 2). Such line widths are not seen in star-forming galaxies but routinely observed in the broad-line regions of AGN, where gas clouds orbit the supermassive black hole [15]. An additional fit to the $[\text{O III}] \lambda 5008 \text{\AA}$ line clearly shows that it is narrow and un-resolved as opposed to the broad Balmer lines. Assuming an SMC extinction law [16], the Balmer line ratio ($H\alpha$ to $H\beta$) indicates a strong dust attenuation of $A_V = 3.0 \pm 0.5$. Note that flatter attenuation curves yield higher A_V . The *Atacama Large Millimeter/sub-millimeter Array* (ALMA) 1.2 mm observations of A2744 [17] firmly rule out dusty star-formation powering the Balmer lines. If caused by star-formation, the emission lines would imply an on-going star-formation rate (SFR) of $\psi \sim 40 \frac{\text{M}_\odot}{\text{yr}}$ and a corresponding 1.2 mm flux of $\sim 1 \text{ mJy}$, whereas no emission is observed to $< 0.1 \text{ mJy}$ (at 3σ). We can therefore clearly rule out the possibility of our source being a star-forming galaxy. However, we cannot entirely rule out that the observed rest-frame UV emission is at least partly of stellar origin. In that case, assuming the entire UV emission comes from stars, we can estimate the stellar mass by assuming a constant star-formation history (SFH) and a formation redshift of $z = 12$ and obtain $M_\star \sim 10^{8.2} \text{ M}_\odot$.

From the width of the $H\beta$ line (see Fig. 2), we derive the black hole mass using standard scaling relations [18], and obtain $M_{\text{BH}} = 4^{+2}_{-1} \times 10^7 \text{ M}_\odot$. We further derive a bolometric luminosity of $L_{\text{bol}} = (1.1 \pm 0.3) \times 10^{45} \frac{\text{erg}}{\text{s}}$ (see Fig. 3) based on the emission line luminosity. Both M_{BH} and L_{bol} are corrected for magnification. The black hole must accrete at a rate of $\dot{M} \sim 0.1 \frac{\text{M}_\odot}{\text{yr}}$ to maintain this luminosity, which roughly corresponds to 30 % of the Eddington limit – set by the balance between gravity and radiation pressure assuming spherical accretion ($L_{\text{bol}}/L_{\text{Edd}} \sim 0.3$).

We can also place a limit on the hidden stellar mass of the host galaxy based on the light profile. All three images are point-sources, despite lensing magnifications differing by a factor 2, which suggests that the intrinsic size of the source is smaller than the point spread function (PSF) of the highest-magnification image. To derive formal upper limits on the size, we fit a Sérsic [19] model to the source in the (highest-resolution) F115W-band, finding the source to be unresolved with a size $r_e < 30 \text{ pc}$ (95 % upper limit) after correcting for the lensing shear (equivalent to less than half the size of the PSF). Adopting an upper stellar density limit equal to that of the densest star clusters [20] or the densest elliptical galaxy progenitors [21] known to date, $\Sigma_\star \sim 5 \times 10^5 \frac{\text{M}_\odot}{\text{pc}^2}$, we can derive an upper limit for how much stellar mass may be contained within $r_e < 30 \text{ pc}$, thus obtaining $M_\star < 1.4 \times 10^9 \text{ M}_\odot$. Our prior photometric fit including ALMA 1.1 mm constraints returns an intrinsic stellar mass upper limit that is consistent with this estimate [3, 4]. The implied $M_{\text{BH}} - M_\star$ -ratio of $> 3 \%$ is extreme compared to local values ($M_{\text{BH}}/M_\star \sim 0.1 \%$) but similar to some other recently discovered high-redshift black holes [22].

The spectrum of A2744-QSO1 (Fig. 2) exhibits some distinct characteristics compared to typical AGN [23] and other moderate-luminosity $z > 4$ AGN discovered with JWST [24]. The optical continuum is very red, and the metal or high-ionization lines (e.g. the $\text{C IV } \lambda\lambda 1548, 1550 \text{\AA}$, $\text{C III} \lambda\lambda 1907, 1909 \text{\AA}$, $\text{Mg II } \lambda\lambda 2796, 2803 \text{\AA}$ and the $[\text{O III}] \lambda\lambda 4959, 5007 \text{\AA}$ -doublets) are very weak compared to the Hydrogen lines (see

Fig. 2 and Tab. 1). Furthermore, using the latest deep X-ray imaging obtained by the *Chandra X-ray Observatory*, we find that all three images are undetected in the X-rays. The combined 3σ upper limit on the rest-frame 40 keV X-ray luminosity of our source is $L_{X,40\text{ keV}} < 3 \times 10^{43} \text{ erg s}^{-1}$, at least $10\times$ weaker than expected based on the UV/optical. This intrinsic X-ray faintness of A2744-QSO1 might suggest a powerful dusty wind, as are seen in low-mass rapidly accreting black holes at low redshift [25]. Such dusty wind-dominated sources are rare at cosmic noon [26], however, and we do not detect any direct evidence for a strong wind in the current data. On the other hand, many black holes radiating at or above their Eddington limits display these unique properties [27], perhaps because of changes in the structure of their accretion disk [28]. Our estimated Eddington ratio of 0.3 would also be consistent with these characteristics.

A2744-QSO1 represents one of the first examples of a heavily reddened broad-line AGN (“red dots”) at high redshift [4], although hints existed also in prior work [29]. These objects are too UV-faint to have been discovered by past ground-based photometric searches [30], but many similar compact red objects have been detected in JWST surveys since [2, 3, 24], as well as other high-redshift AGN [1, 24, 31–33], including behind A2744 [34]. In fact, the UNCOVER “red dot” sample [3] contains a second object at $z_{\text{spec}} = 7.04$ [35], consistent with A2744-QSO1, suggesting these objects may be clustered. Indeed, three out of seven similar objects found in a blank field study [2] were also at similar redshifts to each other. Nevertheless, A2744-QSO1 remains a unique case: As the highest-redshift quasar-like object multiply imaged by a galaxy cluster known to date, it affords exquisite signal-to-noise for such a low-luminosity object.

We derive the spatial number density of A2744-QSO1-type objects by taking the sample of “red dots” detected in UNCOVER [3, 35] and verifying how many of them are confirmed at the same redshift and UV luminosity with the UNCOVER spectroscopy and updated SL model. This yields a number density of $\phi = (7.3 \pm 1.7) \times 10^{-5} \text{ Mpc}^{-3} \text{ mag}^{-1}$, which is consistent with what we found previously based on the UNCOVER photometry [3] ($\phi = 2.5^{+2.6}_{-2.1} \times 10^{-4} \text{ Mpc}^{-3} \text{ mag}^{-1}$) and also with what has been found for reddened $z > 5$ AGN in JWST blank fields [2]. While this represents only a few percent of galaxies at this epoch [36], it is ~ 100 times more numerous than the faintest UV-selected AGN [30]. Because the UV-luminosity of A2744-QSO1 is only $M_{\text{UV},1450} = -16.98 \pm 0.09$, finding this extremely red source at high redshifts required the unprecedented infrared-sensitivity of JWST combined with gravitational lensing. As can be seen in the upper panels of Fig. 3, while A2744-QSO1 is very faint in rest-frame UV compared to other high-redshift AGN, it is nevertheless typical in terms of bolometric luminosity.

A2744-QSO1 is a significant outlier, both in terms of its $M_{\text{BH}} - M_{\star}$ -ratio (which is a lower limit as we do not detect the host galaxy) and the expected number density when compared to theoretical models such as the semi-analytic DELPHI model [37–39]. Assuming black hole seeds with a seed mass of $150 M_{\odot}$ originating from metal-free stars, the DELPHI model predicts that black holes as massive as A2744-QSO1 should be hosted by galaxies with stellar masses of about $M_{\star} \simeq 10^{10.5} M_{\odot}$, i.e. much lower

$M_{\text{BH}} - M_*$ ratios (see Fig. 3), and it under-predicts the number density of A2744-QSO1-like objects. Hydrodynamic simulations make similar predictions at $z \sim 6$ [39], although typically starting with heavier seeds ($10^{4-6} M_{\odot}$). In general, the models fail to predict the high number density associated with A2744-QSO1, which can only be reproduced by the most extreme models without feedback (Fig. 3). This is in conflict with the hints of strong outflows in our source discussed above.

The discovery of A2744-QSO1 shows that our understanding of black hole growth in the early Universe is still incomplete. The photometric samples suggest high number densities, such that a few percent of galaxies at $z > 7$ harbor reddened and rapidly growing black holes. Deep spectroscopic follow-up with JWST will be needed to confirm these number densities and elucidate these objects' properties. The hints of extreme $M_{\text{BH}} - M_*$ -ratios, of a few percent or more, will put strong new constraints on when black holes form relative to their host galaxies. Black holes that outgrow their galaxies put strain on AGN evolution models and might hint at significant dark matter accretion onto primordial seeds [40] or at hyper-compact star-bursting clusters [41] in the early Universe. There are many outstanding puzzles about these objects remaining that JWST is only beginning to address. For A2744-QSO1 in particular, higher-resolution spectroscopy in the rest-frame UV could reveal the possible presence of a dusty outflow in this source, while observations at longer wavelengths could help to better constrain the mass of its host galaxy.

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Table 1 Emission lines detected in the combined spectrum of A2744-QSO1. The stacked spectrum is presented in Fig. 2. Columns represent, from left to right: Emission line designation, rest-frame wavelength, observed wavelength, measured line flux (not corrected for magnification), rest-frame equivalent width (if measured). Upper limits are at 3σ . For Lyman- α , the continuum is based on the extrapolated rest-frame UV continuum assuming a power-law.

Emission Line	λ_0 [Å]	λ_{obs} [μm]	μF_{line} [$10^{-20} \frac{\text{erg}}{\text{s cm}^2}$]	EW ₀ [Å]
Ly α	1216	0.978	236.8 ± 18.8	65.5 ± 12.1
Mg II	2796, 2803	2.252	< 11	-
He I	3889	3.129	18.3 ± 3.3	-
[Ne III]	3968	3.192	15.2 ± 2.5	-
H δ	4102	3.300	19.0 ± 2.4	-
H γ^{a}	4341	3.492	34.5 ± 3.4	-
H β	4863	3.912	100.8 ± 3.3	42.5 ± 2.1
[O III]	4960	3.990	5.0 ± 2.6	-
[O III]	5008	4.029	15.5 ± 2.6	-
Si II ^b	6347	5.106	21.3 ± 3.2	-
[N II]	6550	5.270	76.9 ± 19.1	-
H α	6565	5.282	753.2 ± 29.1	158.0 ± 11

^a Blended with [O III] λ 4363Å.

^b Blended with [O I] λ 6364Å and [Fe x] λ 6375Å.

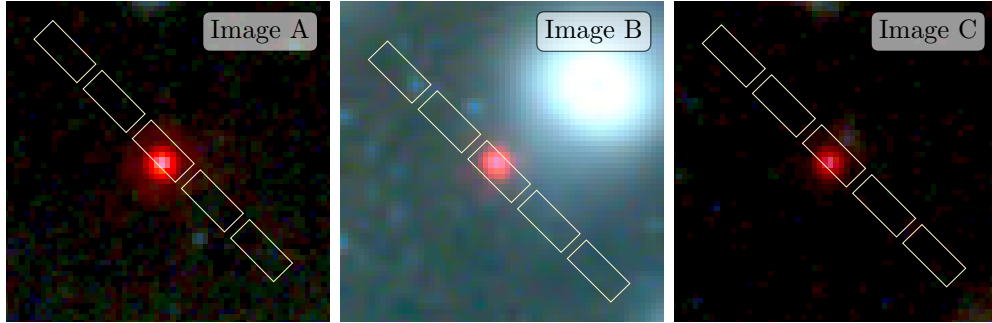


Fig. 1 Composite-color image cutouts of the three images of A2744-QSO1. The $2.4'' \times 2.4''$ cutouts of the UNCOVER JWST-NIRCam composite-color image [12] (Blue: F115W+F150W, Green: F200W+F277W, Red: F2356W+F410M+F444W) show the three images of A2744-QSO1 overlaid with the NIRSpect MSA slitlets of one of our MSA masks as an example (the three images were targeted with several masks and the exact position of the object in each slit may slightly shift).

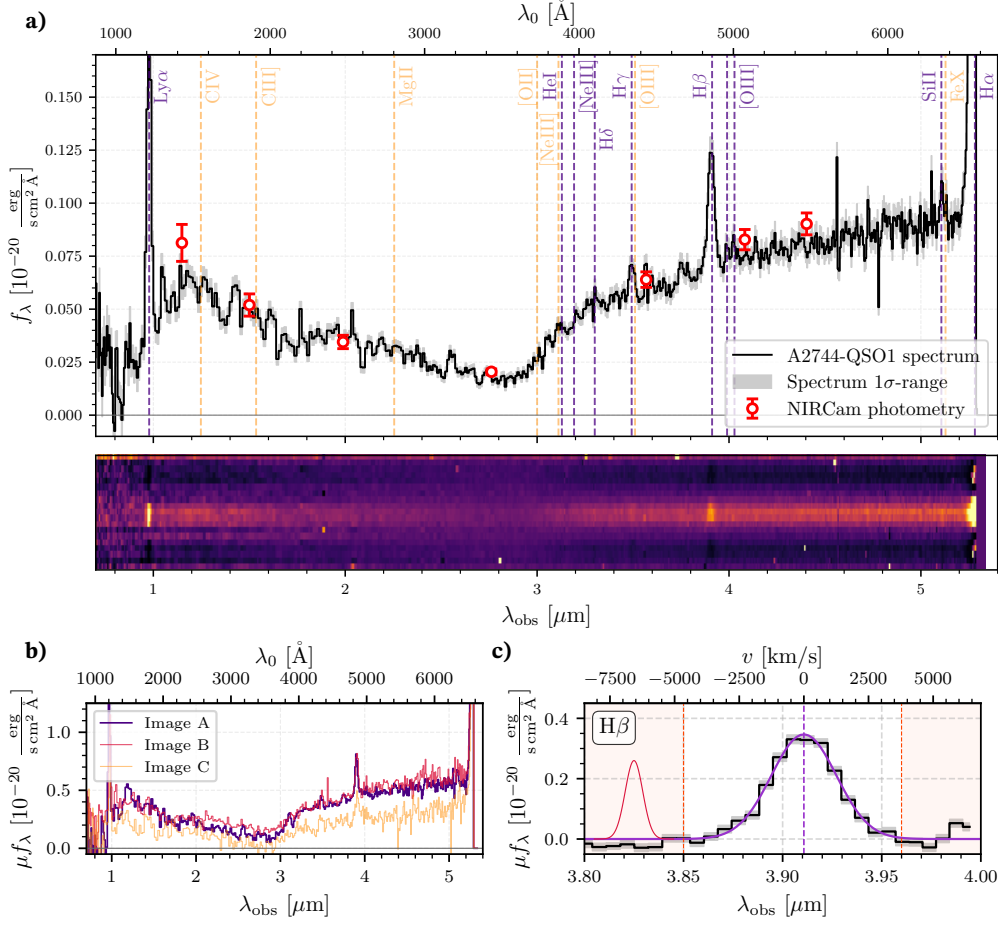


Fig. 2 NIRSpec-prism spectrum of A2744-QSO1. Panel a) shows the de-magnified stacked spectrum (black), equivalent to ~ 1700 h integration time on target once the lensing is factored-in, and its 1σ -uncertainty range (grey). In addition to its red rest-frame optical continuum [4], the source clearly displays strong Hydrogen emission lines in the Balmer-series and Lyman- α . The emission lines allow us to precisely measure the spectroscopic redshift at $z_{\text{spec}} = 7.0451 \pm 0.0005$. We overlay the NIRCам photometry measured previously [4] as red circles. Securely identified and detected ($\geq 3\sigma$) emission lines are marked in purple and selected non-detected lines are shown in orange. Panel b) shows the individual (magnified) spectra of the three images. All three images perfectly align in wavelength space, thus confirming the triply-imaged nature of our object. Panel c) shows the continuum-subtracted H β line (black) and our Gaussian fit to the line (purple). The red shaded areas delimit the region in which the fit is performed and the dark red curve illustrates the NIRCам-prism LSF at the wavelength of H β . All errors shown represent 1σ uncertainties.

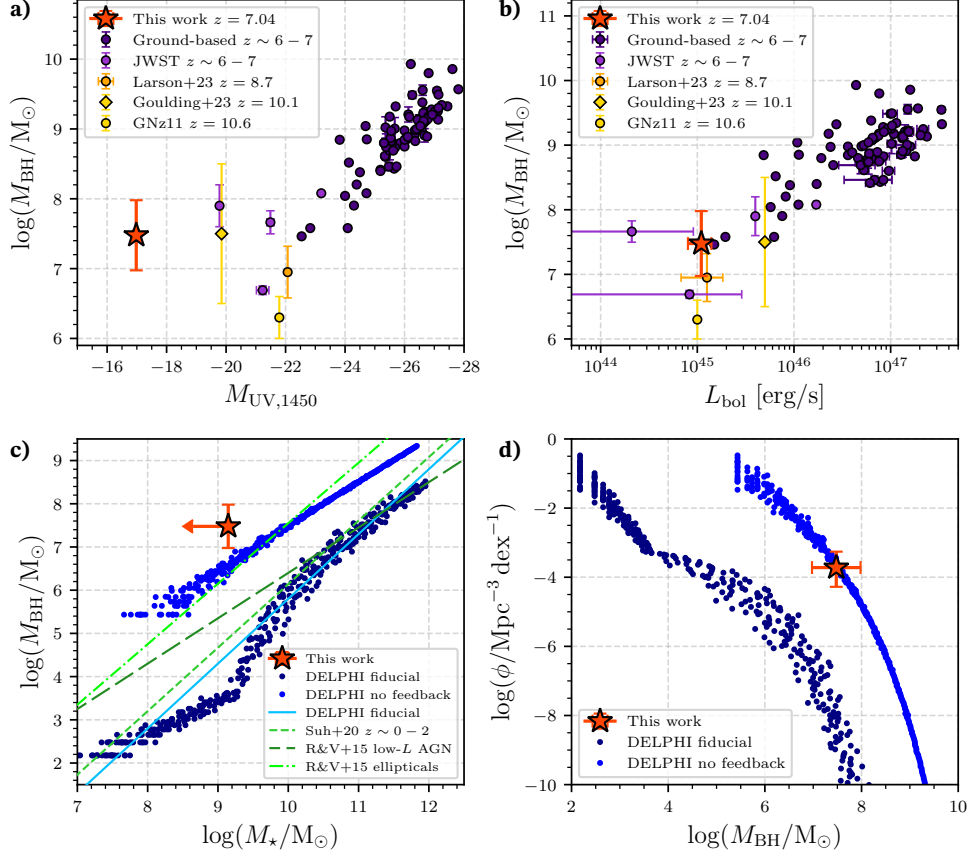


Fig. 3 A2744-QSO1 compared to other broad-line AGN. The top row shows A2744-QSO1 (red star) in the black hole mass-luminosity space of high-redshift broad-line AGN. Purple circles show UV-bright AGN observed from the ground [22, 42], the violet circles show AGN detected with JWST at $z \sim 6-7$ [24, 29, 33], and the orange and yellow symbols show recently confirmed AGN at higher redshifts [31, 32, 34]. While A2744-QSO1 is notably fainter in the UV than other AGN recently detected with JWST (a) because of its strong dust attenuation, it is typical in the $M_{\text{BH}} - L_{\text{bol}}$ plane (b). The bottom row shows A2744-QSO1 in $M_{\text{BH}} - M_{\star}$ -space (c) and in number-densities (d) compared to the DELPHI simulations [37, 38] (blue) and AGN observed at low redshift [43, 44] (green). The dark blue points show the default DELPHI runs, which predict much higher host stellar masses and much lower number density for a black hole of this mass. The light blue points show the models with feedback turned off, in which case the models can approach the observations of A2744-QSO1. Error bars represent 1σ uncertainties.

Methods

Throughout this work, we assume a standard flat Λ CDM cosmology with $H_0 = 70 \frac{\text{km}}{\text{s Mpc}}$, $\Omega_\Lambda = 0.7$, and $\Omega_m = 0.3$. Magnitudes are quoted in the AB system [45] and uncertainties represent 1σ ranges unless stated otherwise.

JWST, ALMA, and *Chandra* observations

The UNCOVER observations used in this work comprise both ultra-deep JWST [46, 47] imaging and spectroscopy of the A2744-field with NIRCam [8] and NIRSpec [14, 48] to 5σ -depths of 29.9 magnitudes each [9, 49]. Since gravitational magnification on average adds ~ 2 magnitudes to the depth in the SL regime, UNCOVER reaches effective depths of ~ 32 AB once the magnification is factored-in. The NIRCam data cover a total area of $\sim 45 \text{ arcmin}^2$ in six broad- and one medium-band filter: F115W, F150W, F200W, F277W, F356W, F410M and F444W. The data were reduced and drizzled into mosaics with the **Grism redshift and line analysis software for space-based slitless spectroscopy** [50] (grizli; v1.6.0). The UNCOVER mosaics, which also include other public JWST imaging data (i.e. the JWST-GLASS imaging [51] and DDT program 2756; PI: W. Chen; as well as archival HST imaging) and were made publicly available on the UNCOVER website. The photometry for all three images of A2744-QSO1 can be found in Table 1 of the discovery paper for A2744-QSO1 [4].

The NIRSpec component of the UNCOVER survey consists of seven micro-shutter array (MSA) [52] pointings on the A2744-field with various exposure times between ~ 3 h and ~ 17 h, totalling 24 h NIRSpec time and 680 targeted objects. The observations follow a 2-point dither pattern and a 3-shutter slit-let nod pattern at a position angle of 266° . The observations were taken from July 31 - August 2, 2023, using the low-resolution prism disperser, which covers the wavelength range $0.7 - 5.3 \mu\text{m}$ at spectral resolutions $R \sim 30 - 300$. The exposure times for the three images of A2744-QSO1 ranged from 9.3 - 16.4 h. The prism data were reduced with MSAEXP [53] (v0.6.10). The code works on individual frames to correct for $1/f$ -noise, find and mask snowballs, and remove the bias. The pipeline first applies a WCS, and then identifies each slit and applies the flat-fielding and photometric corrections. The 2D-slits are then extracted and drizzled onto a common grid. The details of the observational setup and reduced data products, planned for public release before JWST Cycle 3, will be presented in Price et al. (in prep.). All three images of A2744-QSO1 were included in the MSA designs with cumulated exposure times 17.4 h, 9.9 h and 14.7 h for images A, B and C respectively.

In addition, A2744 is part of the *ALMA Frontier Fields* survey [54], which obtained first deep 1.1 mm imaging of the cluster core. These data have now recently been complemented with additional observations covering the entire UNCOVER footprint (Program ID: 2022.1.00073.S; PI: S. Fujimoto), achieving 1σ depths down to $32.7 \mu\text{Jy/beam}$ over an area of $\sim 24 \text{ arcmin}^2$. The UNCOVER ALMA data were also publicly released and described in detail in Fujimoto et al. [17].

A2744 has also been the subject of sensitive X-ray follow-up observations using the *Chandra X-ray observatory*. The 1.25 Ms of *Advanced CCD Imaging Spectrometer*

(ACIS) data used here were previously reduced, analyzed and presented in the supplementary materials section of Bogdán et al. [55]. We use these data products to perform forced photometry at the locations of the three images of A2744-QSO1. Images A and B are heavily contaminated by the foreground X-ray gas associated with A2744 itself. Image C provides the cleanest (lowest-background) region to perform an X-ray analysis of A2744-QSO1. However, we find no significant detection of an X-ray point source. We place a 3σ upper limit of < 12.8 photons in the $2 - 10$ keV energy band, i.e. a flux limit of $f_{X,2-10\text{ keV}} < 2.8 \times 10^{-16} \frac{\text{erg}}{\text{s cm}^2}$ assuming a typical AGN X-ray power-law slope of $\Gamma = 1.9$. This is equivalent to a de-magnified X-ray luminosity of $L_{X,40\text{ keV}} < 5 \times 10^{43} \frac{\text{erg}}{\text{s}}$ at a rest-frame energy of ~ 40 keV given the high redshift of A2744-QSO1. Stacking and de-magnifying all three images produces a marginal increase in sensitivity, owing to the increased magnification but higher backgrounds of images A and B of $L_{X,40\text{ keV}} < 3 \times 10^{43} \frac{\text{erg}}{\text{s}}$.

Gravitational lensing

Throughout this work we use the most recent SL model of A2744 by Furtak et al. [12], which is based on UNCOVER data and also publicly available on the UNCOVER website. The model is slightly updated here with one new multiple image system and an additional spectroscopic redshift [56]. The model was constructed using a new implementation of the Zitrin-parametric code [12, 57, 58], which is not coupled to a fixed grid resolution and thus capable of high resolution results. For the model, the cluster galaxies are modeled each as a double Pseudo Isothermal Ellipsoid [59] (dPIE), scaled according to its luminosity using common scaling relations motivated by the fundamental plane. Three of the brightest cluster galaxies (BCGs), as well as several other galaxies situated close to multiple images, are modeled independently. We use 421 cluster galaxies throughout the UNCOVER $\sim 45 \text{ arcmin}^2$ field-of-view, more than half of which are spectroscopically confirmed [60]. The rest were chosen based on the cluster's red-sequence, identified in a color-magnitude diagram based on the HST/ACS F606W- and F814W-bands bracketing their $4000 \text{ \AA}$ -break. Five cluster-scale dark matter (DM) halos are employed, each modeled as a Pseudo Isothermal Elliptical Mass Distribution [61, 62] (PIEMD). These DM halos are centered on the five BCGs but for three of them, the central position can be iterated for in the minimization process within a couple of arc-seconds around their respective BCG's position. As constraints, 141 multiple images (belonging to 48 sources) are used, 96 of which have spectroscopic redshifts. The vast majority of these is situated in the main cluster core just like A2744-QSO1. The minimization is performed in the source plane using a $\sim 10^5$ -step MCMC process. For the minimization and uncertainty estimation, we use a positional uncertainty of $0.5''$ for each multiple image. The resulting model has a very good image reproduction RMS of $\Delta_{\text{RMS}} = 0.51''$ in the lens plane. We infer magnifications and 95% confidence intervals of $\mu_A = 6.15 [5.76, 6.92]$, $\mu_B = 7.29 [5.11, 7.65]$ and $\mu_C = 3.55 [3.31, 3.80]$ for images A, B and C respectively. These are only very slightly different from the magnifications presented in Furtak et al. [4] and are in generally good agreement with the photometric flux ratios.

Spectroscopic analysis

We use the image multiplicity of A2744-QSO1 to achieve maximum signal-to-noise by stacking the weighted, local-background subtracted and co-added spectra of the three images. The spectrum is then extracted with `MSAEXP` using an optimal extraction [63] and collapsed to a 1D-spectrum. Taking the magnification into account, the effective exposure time of this stacked spectrum is of about 1700 h. The flux-calibration is then performed by bootstrapping to the total photometry of image A, which has the cleanest photometric measurement [4]. Wavelength-dependent slit-losses are modeled with a first-order polynomial for each MSA by convolving with the filter band-pass and solving for the polynomial coefficients by comparing with the photometry. While our stacking and flux-calibration procedures are optimal for maximizing the signal-to-noise and measuring properties such as redshift, black hole mass, etc., they unfortunately smooth out any possible variation between the three images due to the gravitational time delay. Studies of the time-variability of multiply-imaged AGN is important, for example allowing for cosmography studies and reverberation-mapping. This requires careful calibration of the individual spectra though and goes beyond the scope of this work.

The final stacked spectrum of A2744-QSO1 displays numerous emission lines and is shown in Fig. 2. For the most robust emission line in our spectrum, the Balmer line $H\beta$, we use `specutils` [64] (v1.10.0) to fit a two-component (broad and narrow) Gaussian profile to the continuum-subtracted emission line and estimate the uncertainties with an MCMC-analysis using `emcee` [65] (v3.1.4). Our line-fit explicitly includes the NIRSpect point-source line-spread-function (LSF) [66] into the model at the wavelength of $H\beta$ by convolving the model line profile with the LSF in each step of the MCMC before calculating the likelihood. The $\text{FWHM} = 2800 \pm 250 \frac{\text{km}}{\text{s}}$ used in the following is therefore intrinsic, i.e. corrected for instrumental effects. $H\alpha$ is more challenging to model as the red wing of the spectrum falls off the detector and is blended with $[\text{N II}] \lambda\lambda 6550, 6585 \text{\AA}$. In addition, the spectral resolution appears to be high enough at these wavelengths to resolve the narrow component on top of the broad wing. For completeness, we therefore proceed to model $H\alpha$, $[\text{N II}] \lambda\lambda 6550$, $H\beta$, and $[\text{O III}]$ jointly, fixing the metal lines to $100 \frac{\text{km}}{\text{s}}$ FWHM, restricting the narrow-to-broad line ratio to be the same for $H\beta$ and $H\alpha$, but allowing for different line widths between $H\beta$, and $H\alpha$. Encouragingly, the fit parameters are well constrained and reproduce the broad line width of $H\beta$, while finding a smaller FWHM for the broad component of $H\alpha$ $\text{FWHM} \simeq 2300 \pm 250 \frac{\text{km}}{\text{s}}$ as is commonly seen in low-redshift AGN [18]. For consistency, we estimate the width of the narrow lines by fitting the $[\text{O III}] \lambda 5008 \text{\AA}$ line with the same method as above and obtain an LSF-corrected line width that is well below the width of the LSF. This shows that the forbidden lines are indeed narrow. Other spectral lines are identified and modeled with single Gaussians and presented in Tab. 1. Exceptions to this are $H\gamma$, which is blended with $[\text{O III}] \lambda 4363 \text{\AA}$, and $\text{Si II} \lambda 6347 \text{\AA}$, which is blended with $[\text{O I}] \lambda 6364 \text{\AA}$ and $[\text{Fe X}] \lambda 6375 \text{\AA}$. These are modeled similarly to $H\alpha$, i.e. including the blended emission lines, the LSF and the fixed narrow line width. Note that $H\gamma$ also has a resulting a broad FWHM, consistent with $H\alpha$ and $H\beta$.

We derive a Balmer decrement of 7.4 ± 0.4 from the Balmer line measurements in Tab. 1. Assuming an reddening law based on fits to the Small Magellanic Cloud

[16] (SMC) and for $R_V = 2.72$ [67, 68] we find $A_V = 3.0 \pm 0.5$, $A(H\beta) = 3.1 \pm 0.5$, and $A(H\alpha) = 2.1 \pm 0.5$. To date, there have been no direct measurements of the dust attenuation law in obscured high-redshift AGN. The SMC law has however been found to match low-redshift dusty AGN best [69, 70] and also well matches low-metallicity high-redshift galaxies [71–74]. We note that if we assumed a flatter reddening law, the reddening corrections would be larger as would the bolometric luminosity [75]. For example, assuming a Calzetti attenuation law [76], we obtain $A_V = 4.5$. The line fluxes are corrected for dust attenuation before black hole mass and Eddington ratio calculations are made. While the Balmer decrement in the broad-line region arises in part within the broad-line region itself [77], we note that we derive very comparable A_V from the photometric modeling performed in previous work [3, 4].

To investigate a possible SFR origin of the Balmer lines, we convert the dust-corrected $H\alpha$ luminosity to an SFR [78], resulting in an on-going SFR of $\psi \sim 40 \frac{M_\odot}{\text{yr}}$ and predict the rest-frame far-infrared flux using the **Flexible Stellar Population Synthesis** code [79], which implements the Draine & Li [80] dust model. The predicted ALMA 1.2 mm fluxes are $f_{1.2\text{ mm}} \sim 1 \text{ mJy}$ in contrast with a 3σ upper limit of $< 0.1 \text{ mJy}$ based on the deep ALMA map, clearly ruling-out the possibility of star-formation as the origin of the emission lines.

We use the $H\beta$ and $H\alpha$ line widths and attenuation-corrected line fluxes to estimate the single-epoch black hole mass. Because we do not measure the intrinsic continuum, we use the de-reddened and de-magnified Balmer line luminosity to calculate M_{BH} [18]. We find a black hole mass of $M_{\text{BH}} = 4^{+2}_{-1} \times 10^7 M_\odot$, taking the average of $H\alpha$ and $H\beta$, which are consistent with each other. The errors reflect both the line width uncertainties, the differences between the $H\alpha$ and $H\beta$ measurement and the dust attenuation uncertainty. The additional systematic uncertainty term on single-epoch black hole masses is estimated to be $\sim 0.5 \text{ dex}$ [81]. We find a bolometric luminosity of $L_{\text{bol}} = (1.1 \pm 0.3) \times 10^{45} \frac{\text{erg}}{\text{s}}$, assuming a bolometric correction of 10 ± 2 [82], based on the emission line luminosities. For consistency, we also derive the bolometric luminosity from the continuum at rest-frame 5100 Å and find a consistent result. The continuum luminosity estimated from the broad-band photometry in [4] is higher than this estimate by nearly a factor of ten, resulting from uncertainties in quantities like the redshift and the reddening. We thus infer an Eddington ratio of $L_{\text{bol}}/L_{\text{Edd}} \sim 0.3$.

Size measurement

Furtak et al. [4] used **galfit** [83] measurements to the individual images of A2744-QSO1, in all available broad bands from F115W to F444W, together with the fact that all three images appear as point sources despite suffering different amounts of shear and magnification, to conclude that the object was indeed point-like and from the PSF size to obtain a source size estimate of $r_e < 35 \text{ pc}$ [4].

Here we revisit this size estimate with the most up-to-date PSF models and data reduction. Using **pysersic** [84], we fit a single Sérsic model to the F150W image. We find a firm upper limit on the size of $r_e < 11.2^{+1.6}_{-0.8}$ milli-arcseconds, which translates to a 2σ upper limit of $r_e \lesssim 100 \text{ pc}$ for the lensed image. Applying a lensing correction of $\mu_t = 3.52$ – the tangential shear in image A used for the size measurement – we

then confirm a firm 2σ upper limit of 30 pc on the source radius, in agreement with the Furtak et al. measurement [4].

Number density

In order to derive the number density of objects like A2744-QSO1, it is needed to compute the corresponding effective co-moving volume V_{eff} which folds-in the lensing distortion and possible selection effects. This is done following a forward-modeling approach [36], which was re-implemented for JWST observations and in particular for UNCOVER [85]: We use the (de-magnified) SED of A2744-QSO1 to populate the UNCOVER source plane of A2744 [12] with 10000 similar mock sources. These are then deflected into the lens plane and injected in the UNCOVER mosaics on which we re-run the source detection and selection methods used for the catalogs [49, 86]. The ratio of recovered to injected sources allows us to determine the selection-function for objects such as A2744-QSO1. This selection-function is multiplied with the co-moving volume element which is then integrated over the source-plane area of sufficient magnification that allows us to detect the source at 3σ in each NIRCcam band, given the UNCOVER mosaic depths [49]. The thus obtained effective volume for A2744-QSO1 is $V_{\text{eff}} = 27541 \pm 294 \text{ Mpc}^3$.

Given that the $\sim 45 \text{ arcmin}^2$ of the UNCOVER field-of-view [9] contains two “red dot” objects of the redshift and UV luminosity of A2744-QSO1 [3], A2744-QSO1 itself and one other [35] projected 369 kpc away in the source plane, the number density is $\phi = (7.3 \pm 1.7) \times 10^{-5} \text{ Mpc}^{-3} \text{ mag}^{-1}$.

Theoretical expectation

We compare our results to the **Dark matter and the Emergence of gALaxies in the ePoch of reIonization** [37, 38] (DELPHI) semi-analytic model. This is suited to jointly track the assembly of the dark matter, baryonic and black hole components of high-redshift ($z \gtrsim 7$) galaxies up to $z \sim 40$. Starting at $z \sim 7$, we build analytic (binary) merger trees for 700 galaxies equally spaced in log-space in the halo mass range $M_{\text{h}} = 10^8 - 10^{15} \text{ M}_{\odot}$ using a constant time-step of 30 Myr. Each of these halos is assigned a number density according to the Sheth-Tormen halo mass function [87] at $z \sim 7$ and the number density propagated throughout the merger tree.

In this work, we only consider seeding by stellar black hole seeds formed from metal-free (Pop. III) stars. The first halos (starting leaves) of any halo are seeded with a stellar black hole (of $150 \text{ M}_{\odot}$) at $z > 13$ assuming that halos collapsing from high ($\gtrsim 3.5$)- σ fluctuations in the primordial density field are most likely to host such seeds. At any time-step, the initial gas mass is the sum of that brought-in by mergers and smooth accretion. A fraction of this gas can form stars with an efficiency that is the minimum between the Supernovae (SN II) energy required to unbind the rest of the gas and an upper threshold f_* . In terms of black holes, in our *fiducial* model, these grow both through (instantaneous) mergers and accretion. A fraction of the gas mass left in the halo after star-formation and its associated feedback can be accreted onto the black hole with the accretion rate depending on a critical halo mass (which at $z \sim 7$ corresponds to $10^{10.9} \text{ M}_{\odot}$). While black holes in lower mass halos are assumed to

have puffed-up disks, only allowing for low accretion rates ($\sim 7.5 \times 10^{-5}$ Eddington), black holes in halos above this critical mass can accrete a fixed fraction of the gas mass up to the Eddington limit; a fraction of the black hole energy (0.3%) is allowed to couple to the gas content. The *maximal (no feedback)* model presents an upper limit where every black hole is always allowed to accrete the gas mass available, up to the Eddington limit and we ignore gas mass lost in outflows, both due to type II supernova and black hole feedback.

The model has been tuned to reproduce all the key observables for both star-forming galaxies at $z \sim 5-10$ (including the evolving UV luminosity and stellar mass functions) and AGN at $z \sim 5-6$ (including the AGN UV luminosity and black hole mass functions).

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Data availability. The raw JWST data used in this work are publicly available on the **Barbara A. Mikulski Archive for Space Telescopes (MAST)**. These observations specifically are associated with the JWST GO program number 2561, JWST ERS program number 1324, and JWST DD program number 2756. The reduced UNCOVER mosaics, catalogs, lens models as well as spectra are part of the public data release by the UNCOVER team [9, 12, 49] (see also Price et al. in prep.). The public ALMA data can be found on the ESO Science Archive under IDs 2018.1.00035.L and 2013.1.00999.S. The deep *Chandra* image used here comprises data from over 60 individual programs which are available on the *Chandra Data Archive*. The individual observations IDs are listed in Table 1 in Bogdán et al. [55].

Code availability. This research made use of **Astropy**, a community-developed core Python package for Astronomy [88, 89] as well as the packages **NumPy** [90], **SciPy** [91], **Matplotlib** [92] and the **MAAT** Astronomy and Astrophysics tools for **MATLAB** [93].

Competing interests. The authors declare no competing financial interests.

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